

Vorlesung IO:Ausbreitung der kosmischen Strahlung in der Galaxie

Gleichung für Teilchentransport

- Teilchendichte, Diffusion, Quellen
- Energieverlust während des Transports
- Quell- und Verlustterme durch Teilchenerzeugung

Betrachtung von Grenzfällen

- Punktförmige Quelle
- Unendlich ausgedehntes, homogenes Quellgebiet

Vorlesung und Übungen : Daten

Vorlesung: Dienstags

Übungen: Donnerstags

gehalten von Max Stadelmaier

3. Nov. 2020 10. Nov. 2020 17. Nov. 2020 24. Nov. 2020 1. Dez. 2020 8. Dez. 2020 15. Dez. 2020 22. Dez. 2020 12. Jan. 2021 19. Jan. 2021 26. Jan. 2021 2. Feb. 2021 9. Feb. 2021 16. Feb. 2021

 19.11.2020 - Blatt I

 03.12.2020 - Blatt 2

 17.12.2020 - Blatt 3

 14.01.2021 - Blatt 4

 11.02.2021 - Präsentation (Paul Filip)

 18.02.2021 - Blatt 5

Ersatz für 19.1. 4. Feb. 2021

$\langle X_{max} \rangle$ und Streuung



Unsicherheiten durch starke Modellabhängigkeit sehr groß

Phänomenologische Beschreibung:

Energiespektrum und Elementzusammensetzung



Dipolstruktur (Auger)



Figure 1: Map in Equatorial coordinates of the CR flux above 8 EeV, averaged on top-hat windows of 45° radius. The location of the Galactic plane is shown with a dashed line, and the Galactic center is indicated with a star.



Figure 2: Energy dependence of the dipolar amplitude measured above 4 EeV. Also shown are the predictions from scenarios [12] with extragalactic sources.

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Einfluss des galaktischen Magnetfelds auf einen Dipol



Figure 7. Change of the direction of the dipolar component of an extragalactic flux after traversing the Galactic magnetic field, modeled as in Jansson & Farrar (2012). We consider a grid (black circles) corresponding to the directions of a purely dipolar flux outside the Galaxy. Points along the lines indicate the reconstructed directions for different values of the particle rigidity: 32, 16, and 8 EV, and, at the tip of the arrow, 4 EV. The line color indicates the resulting fractional change of the dipole amplitude. The observed direction of the dipole for energies $E \ge 8$ EeV is indicated by the gray plus sign, with the shaded area indicating the 68% CL region. The labels I and O indicate the directions toward the inner and outer spiral arms, respectively.

Dipolentwicklung mit Energie



Figure 1. Reconstructed equatorial dipole amplitude (left) and phase (right). The upper limits at 99% CL are shown for all the energy bins in which the measured amplitude has a chance probability greater than 1%. The gray bands indicate the amplitude and phase for the energy bin $E \ge 8$ EeV. Results from other experiments are shown for comparison (IceCube Collaboration 2012, 2016; KASCADE-Grande Collaboration 2019).

The Imprint of Large Scale Structure on the Ultra-High-Energy Cosmic Ray Sky



The Imprint of Large Scale Structure on the Ultra-High-**Energy Cosmic Ray Sky** 1.10 observation, Auger exposure $I_{max} = 1$ Case d90, Auger exposure $I_{max} = 1$ relative Case d90, full sky $I_{max} \gg 1$ Case d90, full sky, illumination flux Case SH*, Auger exposure $I_{max} = 1$ Observation $E \ge 8 \text{ EeV}$ 45°Top Hat Smoothing Case d90 $E \ge 8$ EeV 45°Top Hat Smoothing Case SH² 0.30 $E \ge 8 \text{ EeV}$ 45°Top Hat Smoothing 0.90 -90 -90 -90 90 1.24 0.20 relative amplitude 0.10 180 flux Observati Case SH² Case d90 $E \ge 32 \text{ EeV}$ $E \ge 32 \text{ EeV}$ $E \ge 32 \text{ EeV}$ dipole ; 0.76 45°Top Hat Smoothing 45°Top Hat Smoothing 45°Top Hat Smoothing 90 90 local Li-Ma 0.05 ο.0 -5.6 16-32 ≥ 32 8-16 ≥ 8 0.03 10 20 30 40 50 Observation Case SH* $E \ge 38 \text{ EeV}$ $E \ge 38 \text{ EeV}$ E [EeV] 27° search radius 27° search radius

Hotspot gesehen durch Teleskoce Array (TA)



Figure 1: (a) A significance map of the UHECR events with E > 57 EeV for 11 years of TA data (May 2008 - May 2019) in the equatorial coordinates. Events are smoothed by 25° oversampling radius circle, which is defined in this paper. (b) A significance map of the UHECR events with E > 57 EeV for events observed in the 1st 5 years of TA data (May 2008 - May 2013). Events are smoothed by 20° oversampling radius circle according to our original paper [4]. The solid curves indicate supergalactic plane (SGP) and the galactic plane (GP).



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AugerPrime Advent of multihybrid data



Ausbreitung der kosmischen Strahlung in der Galaxie

Eigenschaften der kosmischen Strahlung



Vergleich mit Häufigkeit im Sonnensystem





Figure 7: Amplitude (top) and phase (bottom) measurements of the first harmonic in right ascension as a function of energy, from various reports. Amplitudes drawn as triangles with apex pointing down are the most stringent upper limits up to date in the considered energy ranges.

Größenverhältnisse und galaktisches Magnetfeld





Figure 1. Trajectories of iron anti-nuclei backtraced from the Earth. Left panel: Energies equal to $(1, 4, 8) \times 10^{18}$ eV; **Right panel:** Energies equal to $(1, 4, 10) \times 10^{19}$ eV. For details on the Galactic magnetic field model, see text.

(Andromeda, M31)

Magnetfeldkomponenten:

- Reguläre Feldkomponente: Ballistische Bewegung
- Turbulente Feldkomponente: Diffusion





B=3µG=30 nT in der Nähe des Sonnensystems

Mögliche Quellen kosmischer Strahlung



Wechselwirkung mit interstellarem Medium





Transportgleichung: Übersicht



Transportgleichung: Diffusion im Magnetfeld

Exkurs: Ites Ficksches Gesetz (Gastheorie)





Transportgleichung: kontinuierliche Energieverluste



