

Introduction to Cosmology

Winter term 23/24

Lecture 1

Oct. 24, 2023



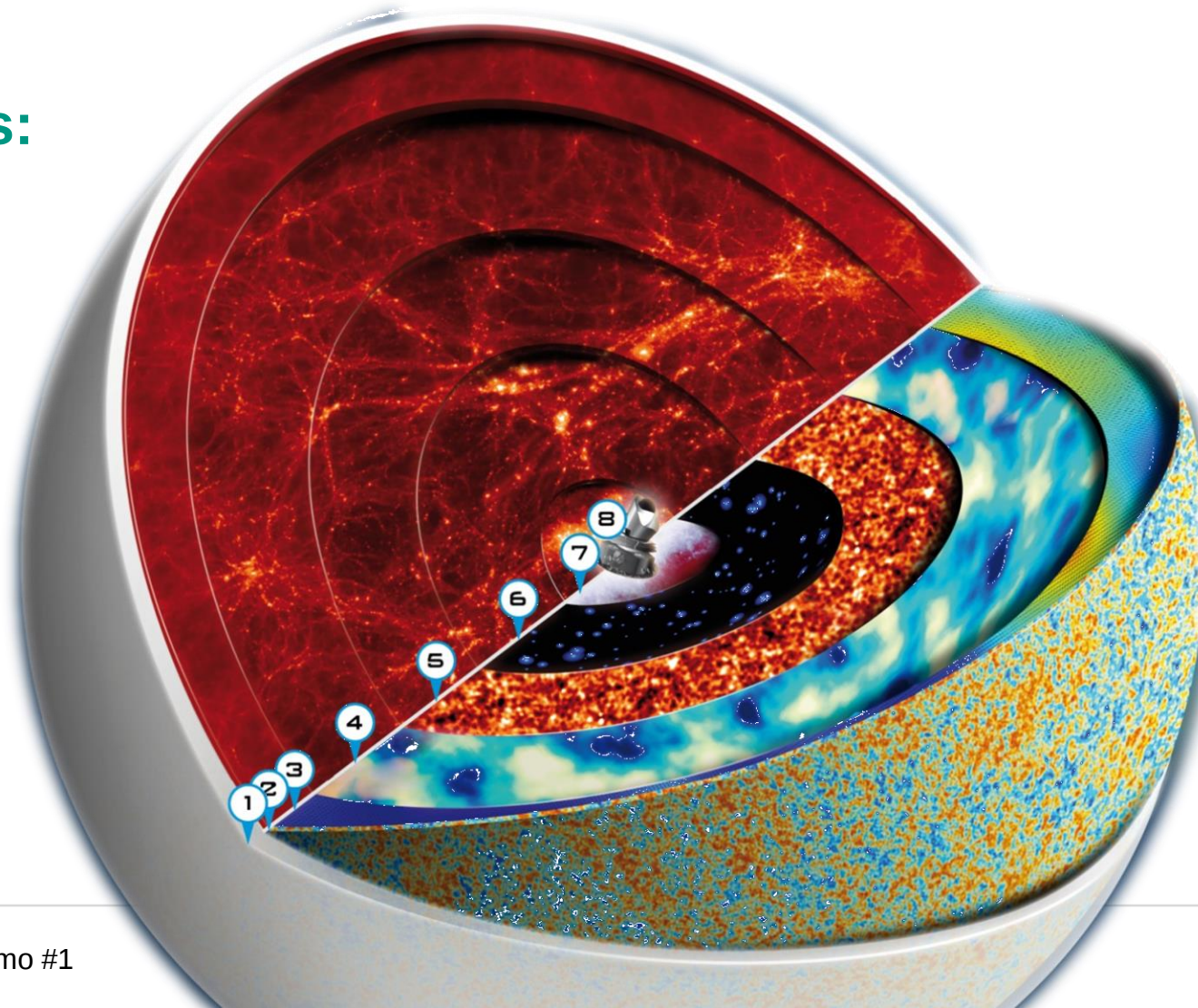
Evolution & Structure of the
Universe from the Big Bang
up to now

Cosmology – Definition

■ Evolution & Structure of the Universe from the **Big Bang** to now

important stopovers:

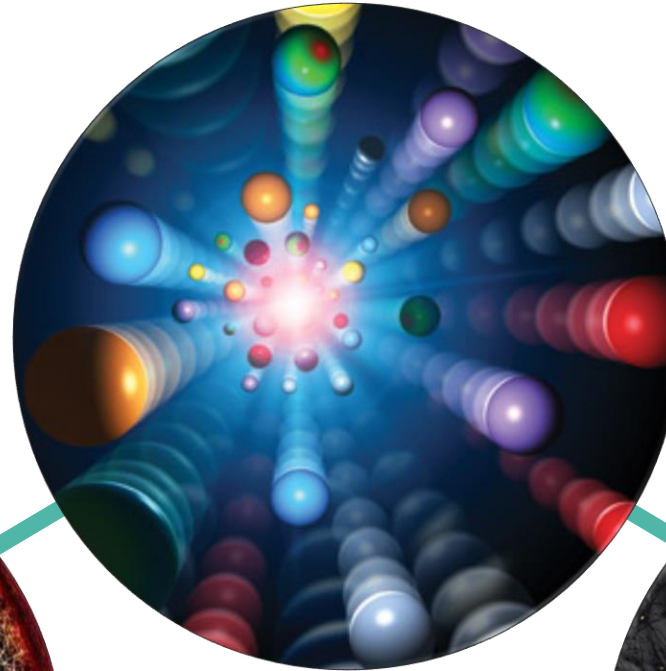
- 1 – **Big Bang**
- 2 – *CMB* (3K)
- 3 – reionisation
- 4 – dark ages
- 5 – first stars
- 6 – first galaxies
- 7 – Milky Way
- 8 – **today**



Cosmology – range of subjects

■ subjects of cosmology

processes in

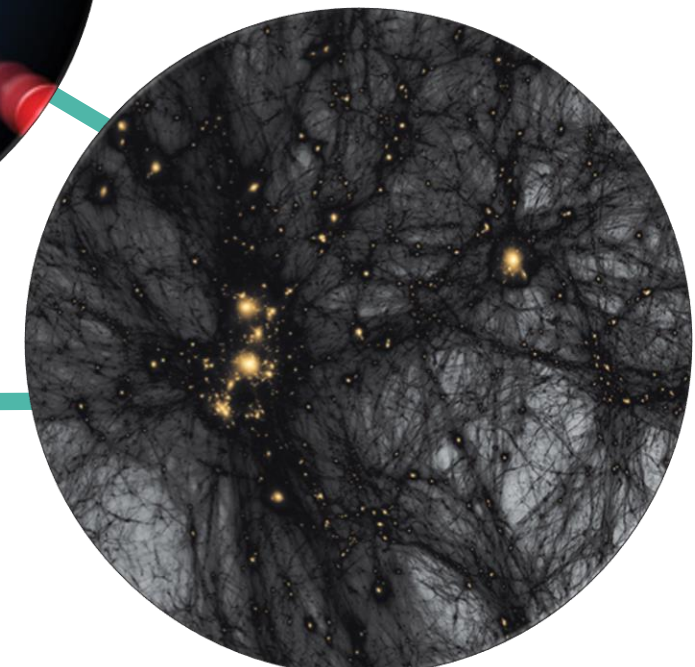


the **Early Universe**

**Structure
Formation**



**three pillars
of these lectures**

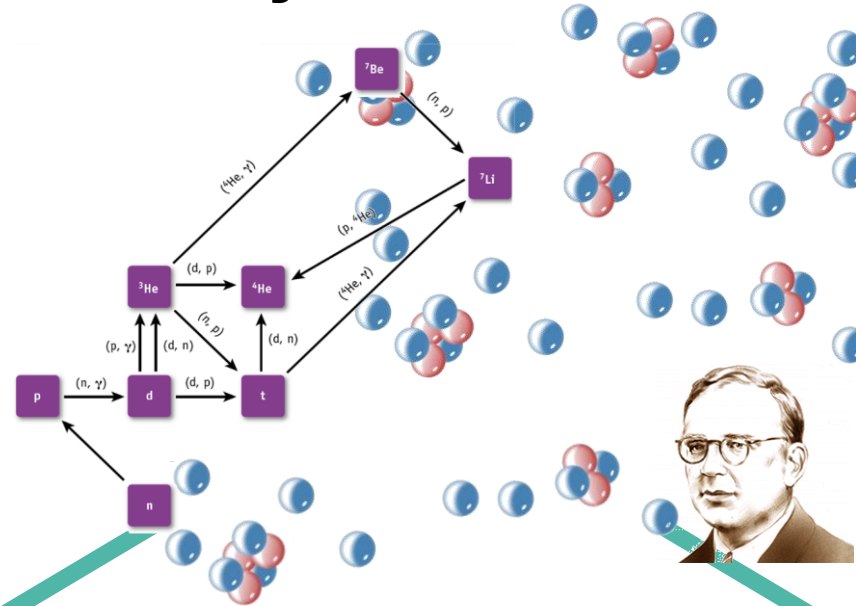
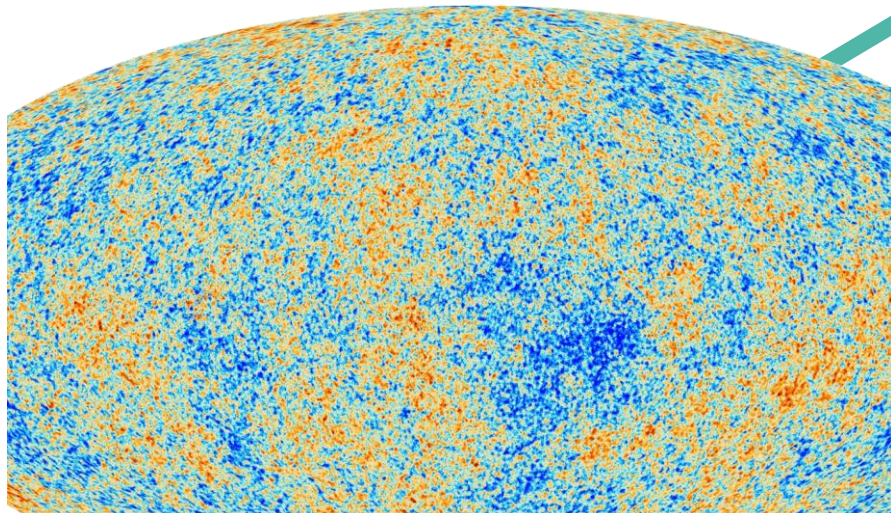


**Dark
Universe**

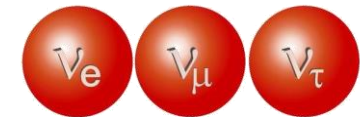
Cosmology – range of subjects

■ example: thermal processes

*CMB** –
temperature



formation of
elements
during Big Bang



neutrino
properties

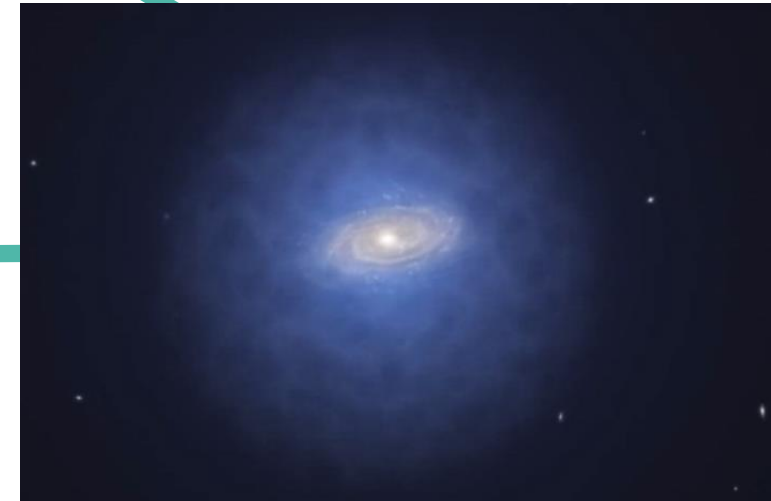
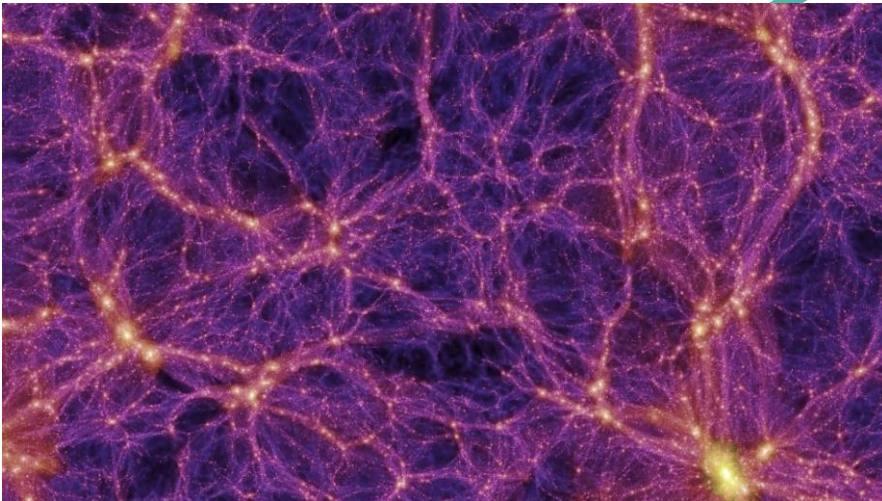
Cosmology – range of subjects

- example:
Dark Matter (DM)

DM in
structure formation



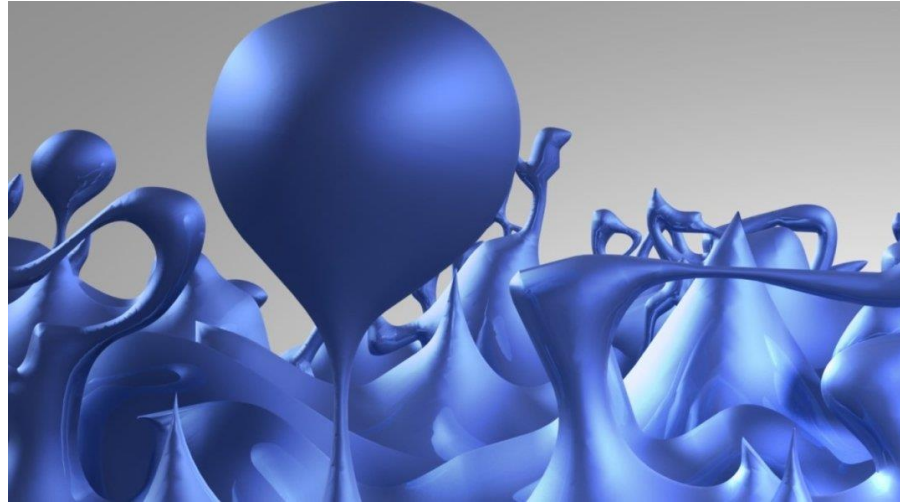
DM – particle candidates:
*HDM**, *WDM***, *CDM****



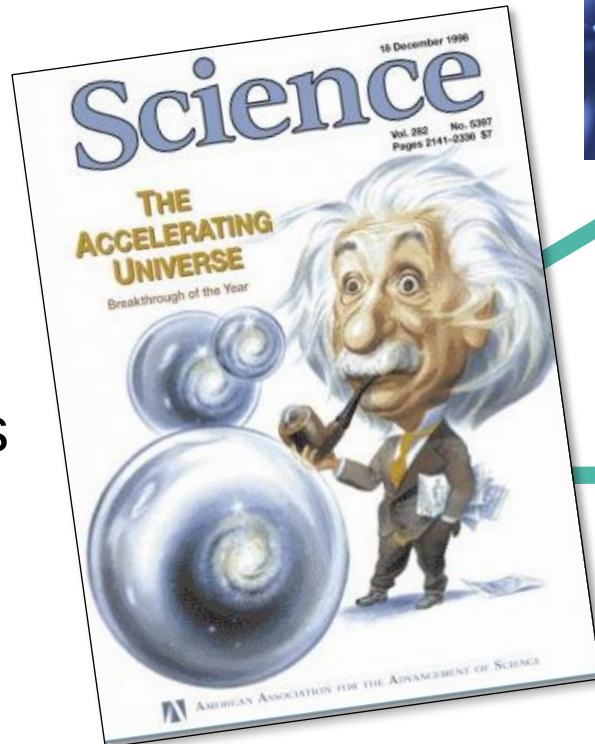
galactic
DM –
halos

Cosmology – range of subjects

- example:
Dark Energy (DE)



nature of **DE**:
**cosmological
constant** ?



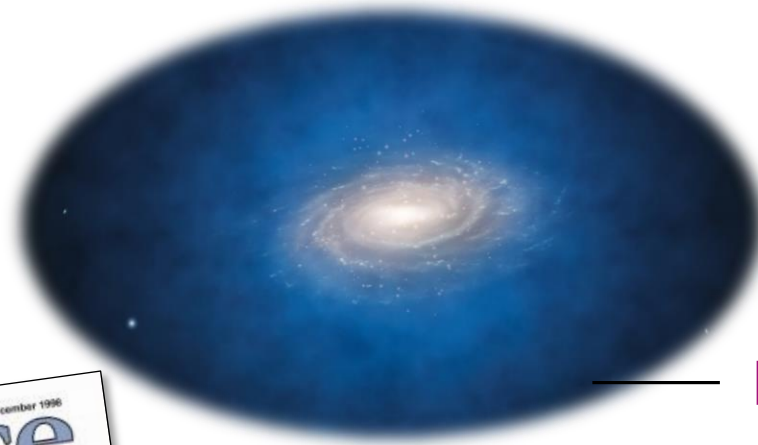
evidences
for **DE**



DE &
expansion
rate of the
universe

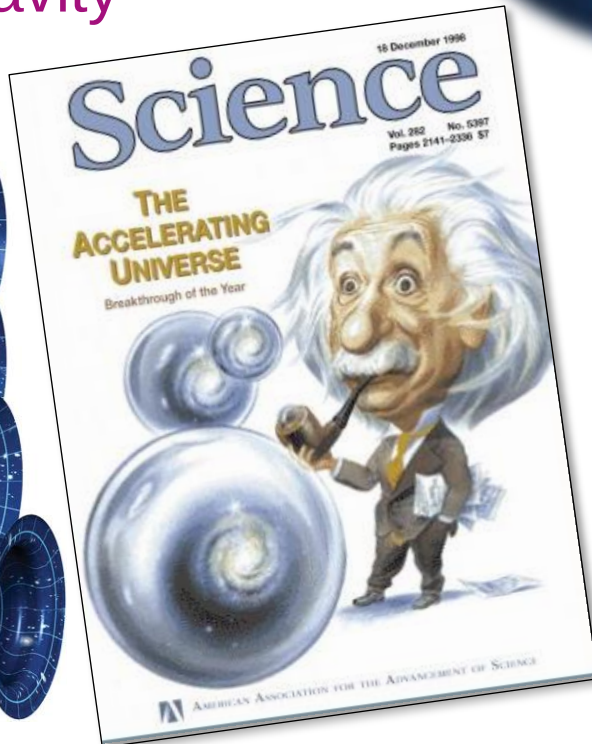
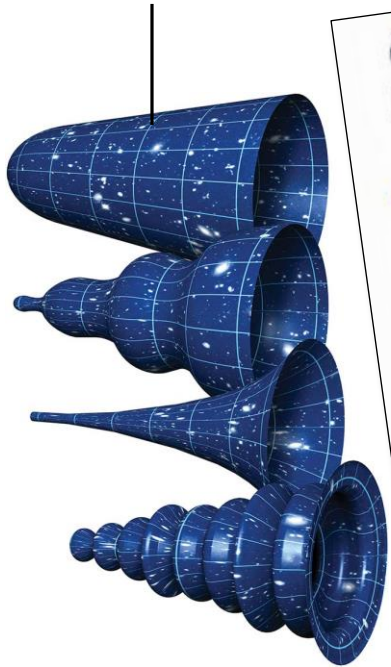
Cosmology – new concepts

■ classification



Dark Matter

Anti-gravity



Bachelor
Studies



gravity

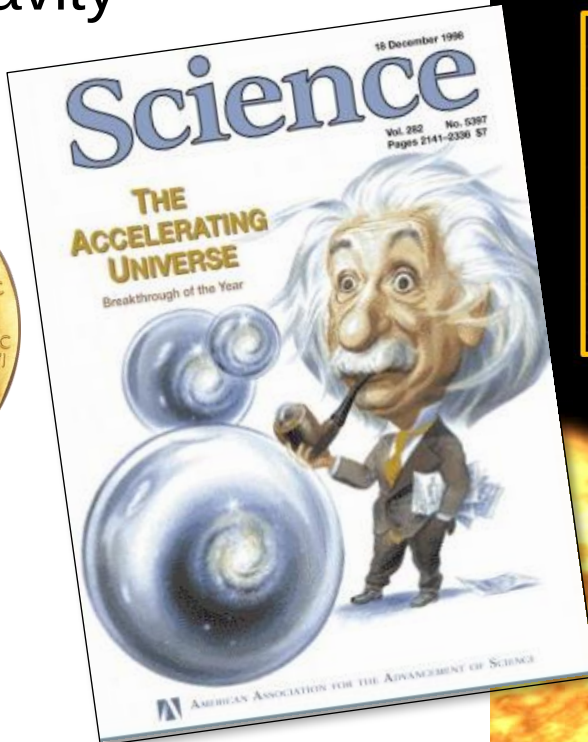
nuclear physics
thermodynamics
optics

condensed
matter, atoms,
molecules

Cosmology – modern observations

■ Nobel Prize 2011

Anti-gravity



„for the discovery of the **accelerating expansion of the Universe** through observations of distant supernovae”



WD

SN Ia



Saul Perlmutter



Adam G. Reiss

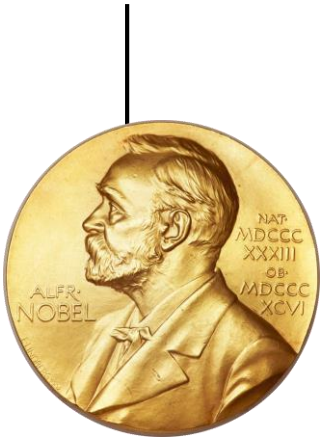


Brian P. Schmidt

Cosmology – modern theories

■ Nobel Prize 2019

structure formation



50 %

$$\delta(\vec{r}) = \frac{\rho(\vec{r}) - \langle \rho \rangle}{\langle \rho \rangle}$$

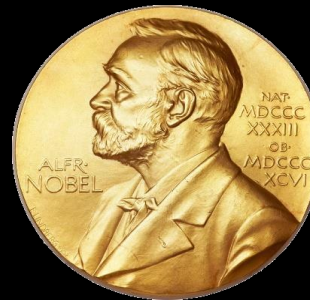
Nobel prize for physics 2019

„for theoretical discoveries
in **physical cosmology**“



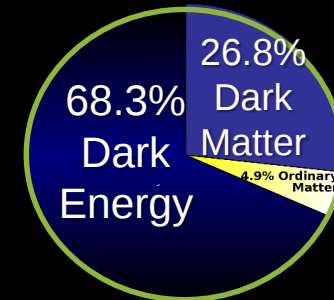
James Peebles

Big
Bang



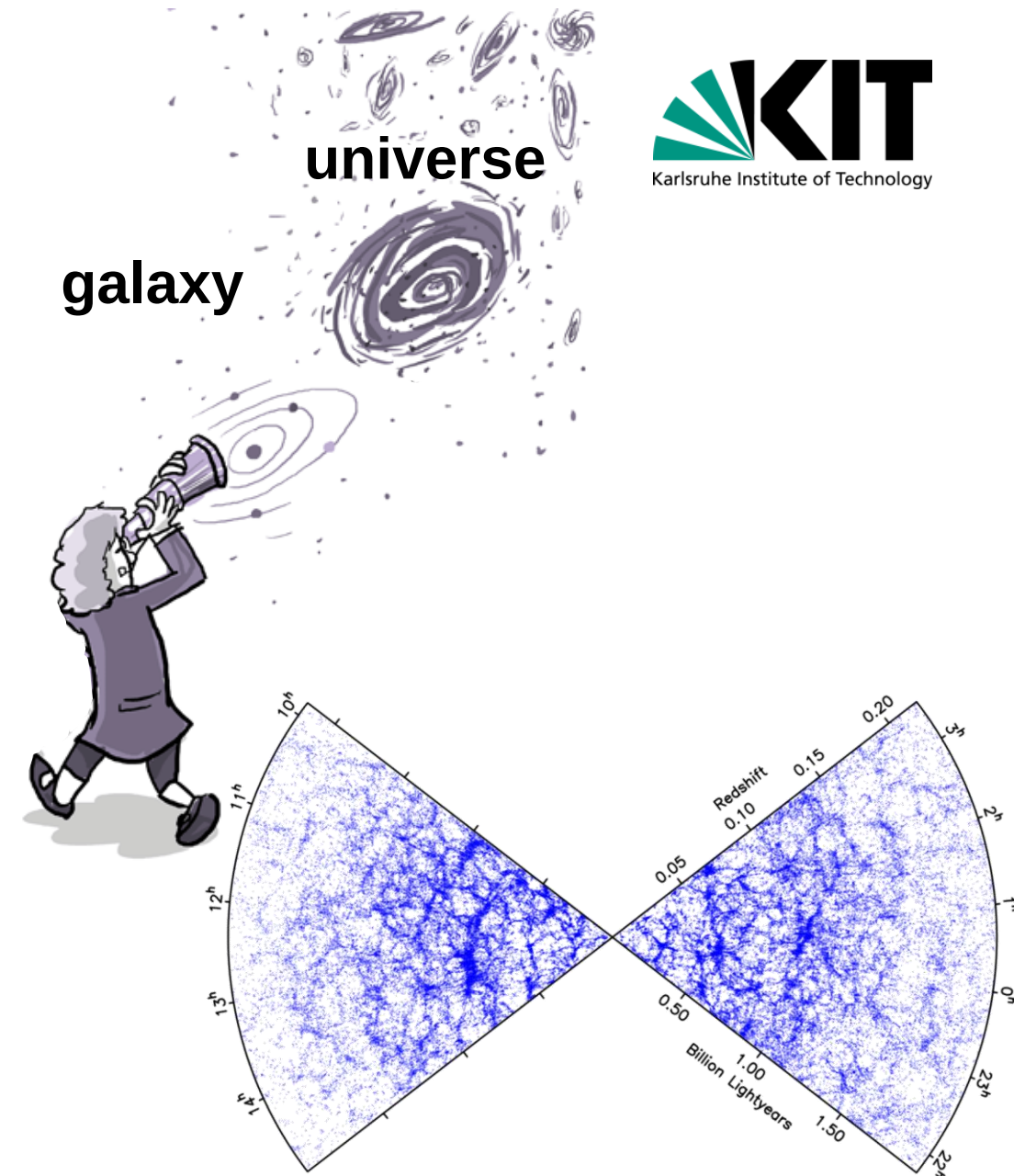
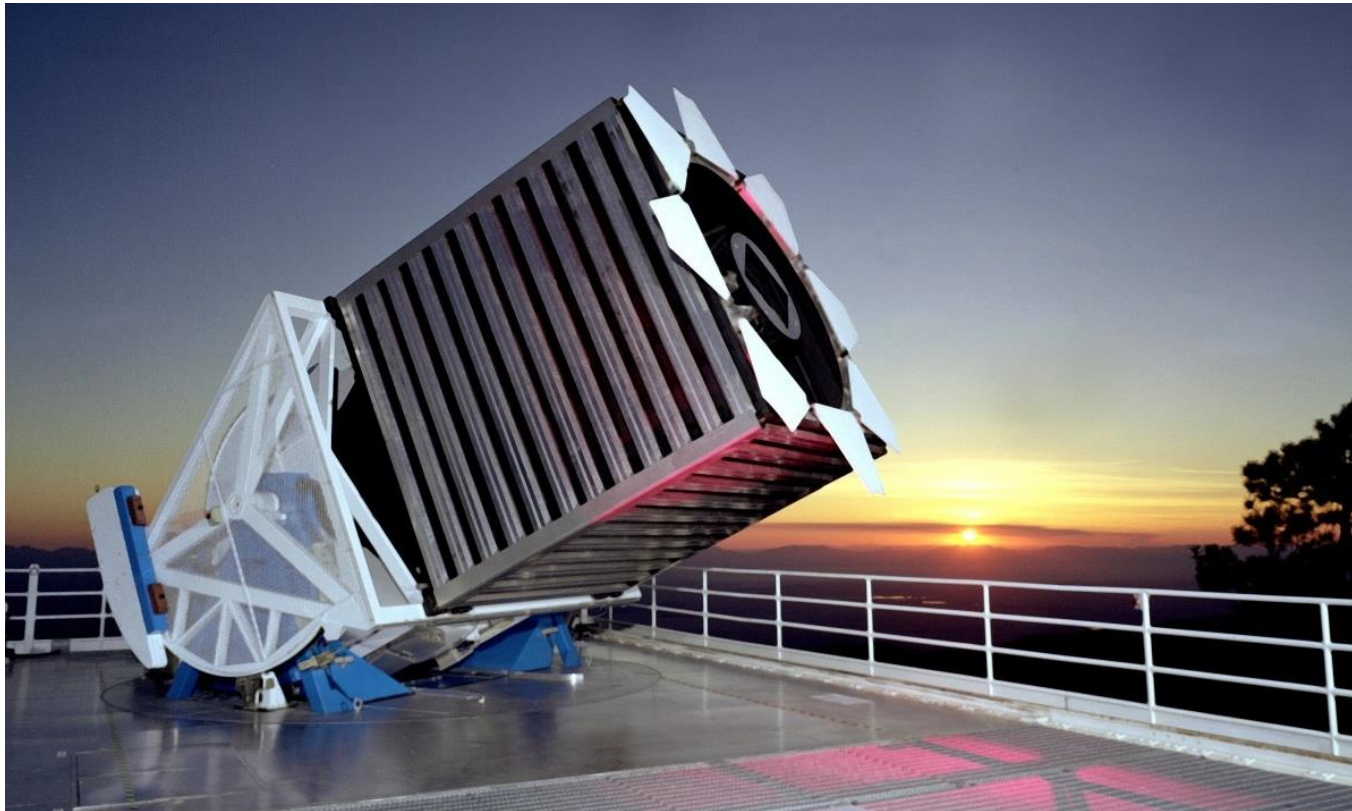
development of large-scale structures

today



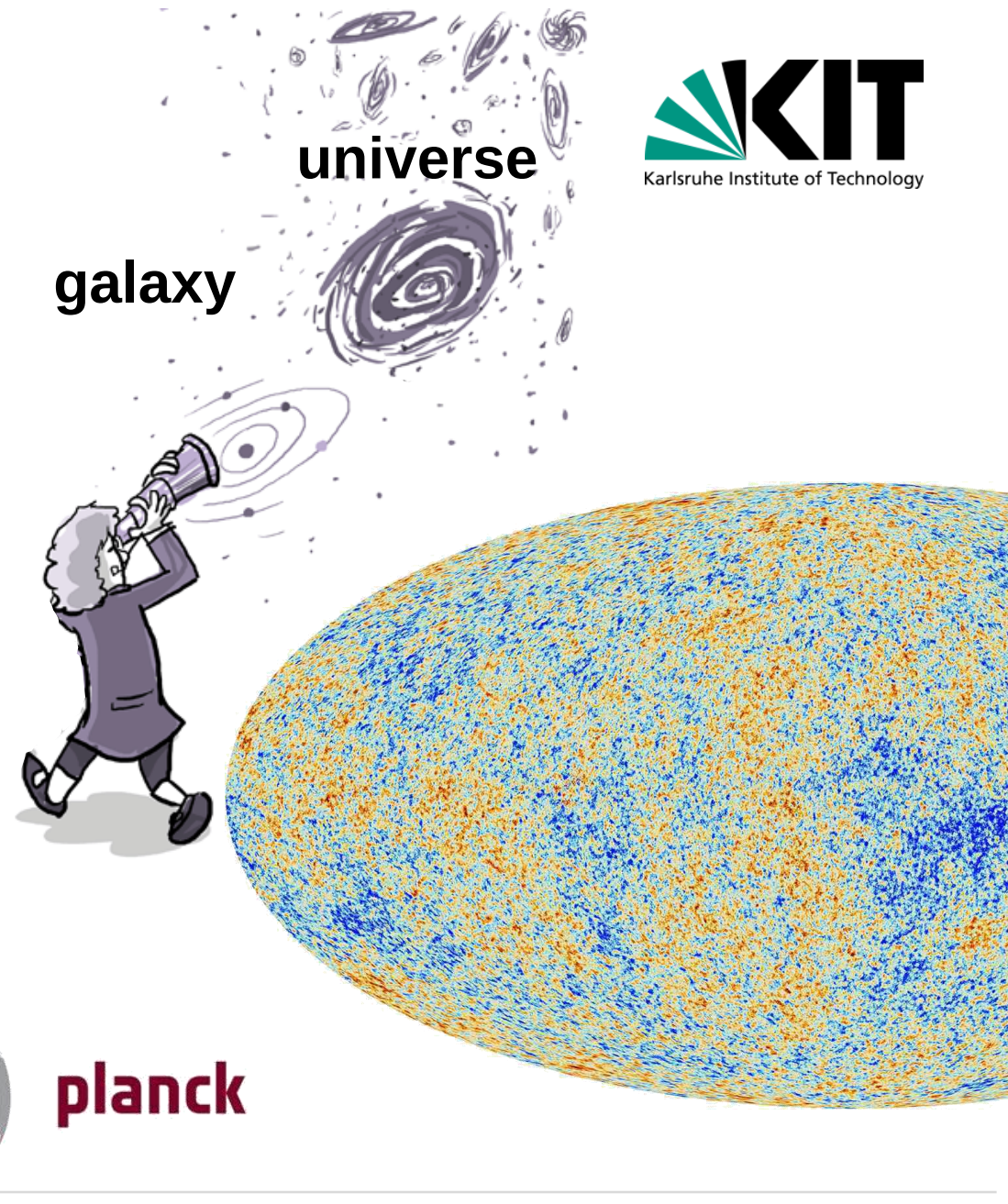
Cosmology – observations

- example: galaxy redshift surveys



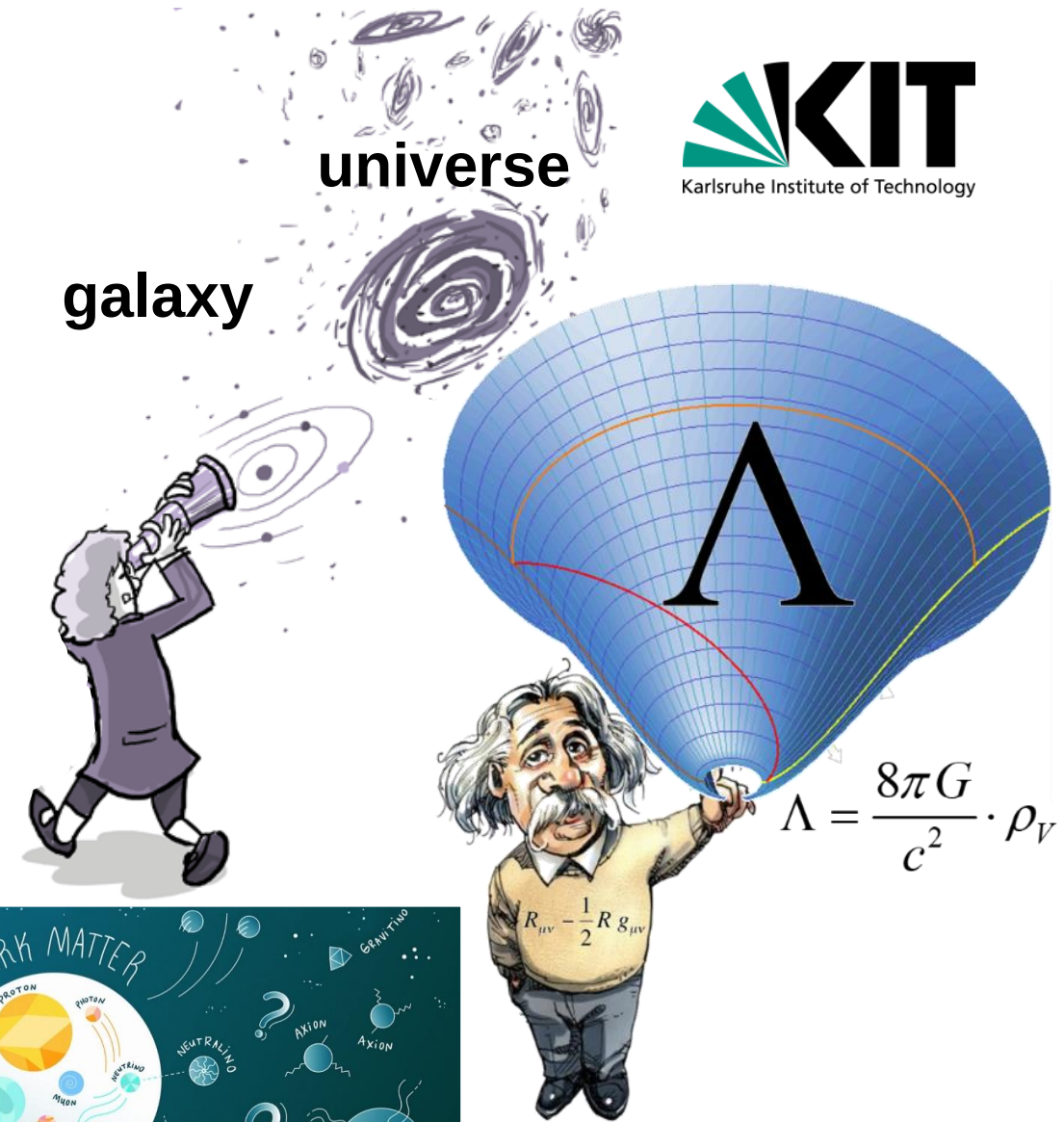
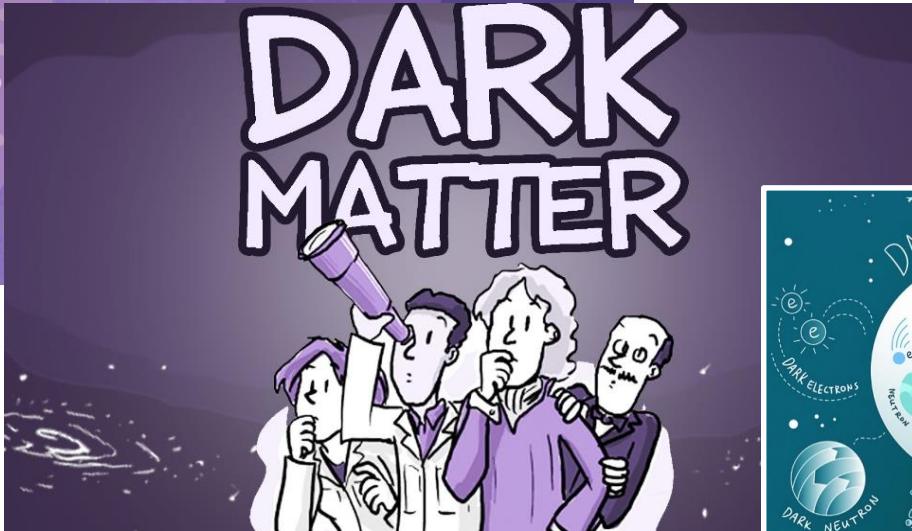
Cosmology – observations

- example: temperature fluctuations of the *CMB* (*C*osmic *M*icrowave *B*ackground radiation)



Cosmology – open questions

- ## ■ what is the intrinsic nature of Dark Matter & Dark Energy ?



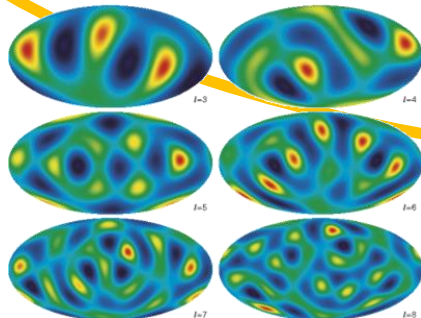
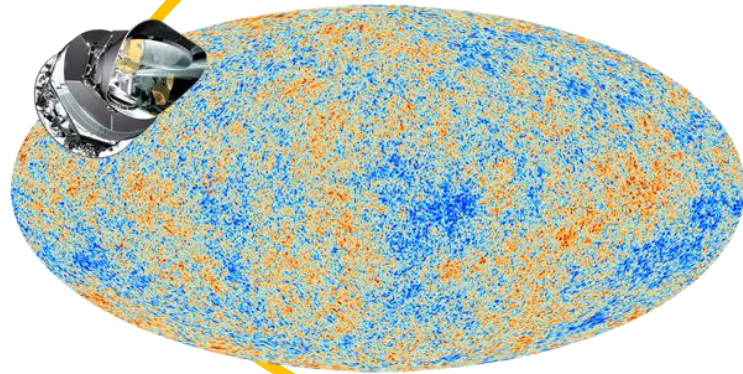
I can learn the following in this lecture series:

**BREAKING
NEWS**

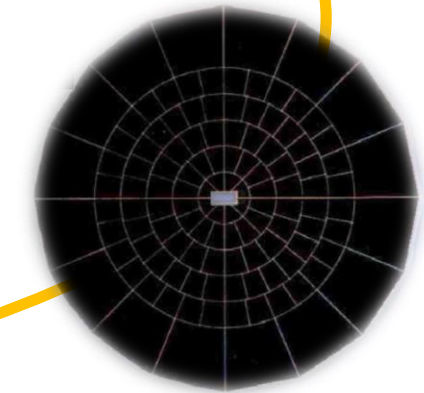


$$\left(\frac{\dot{H}(t)}{H(t)}\right)^2 \sim 1 + \frac{7\Delta N_\nu}{43}$$

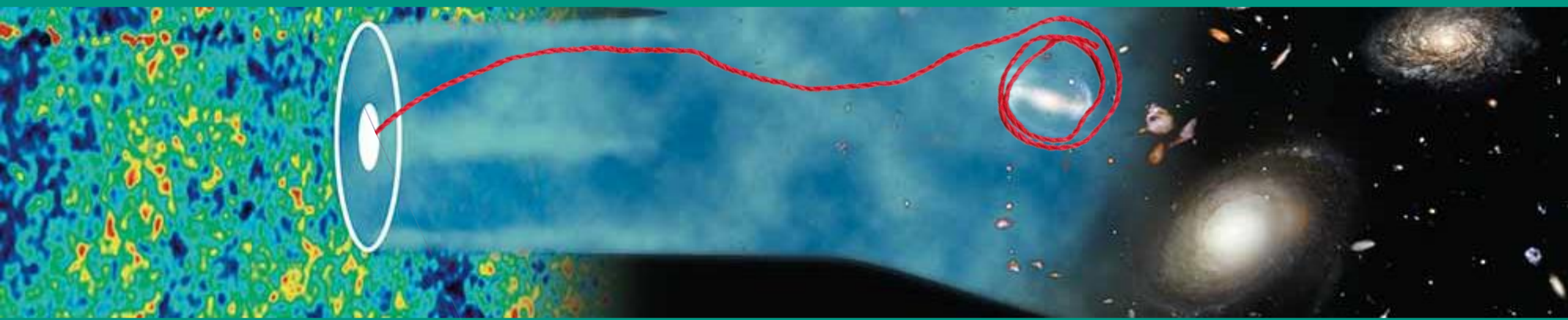
...new tools for
discoveries...



...modern analysis methods...



...advanced experimental
methods...



OUTLINE

1. Introduction

- 1.1 fundamental principles
- 1.2 distance ladders for the Universe

2. Expanding Universe

- 2.1 Hubble Expansion
- 2.2 Friedmann–Lemaître equations

3. Thermal Universe

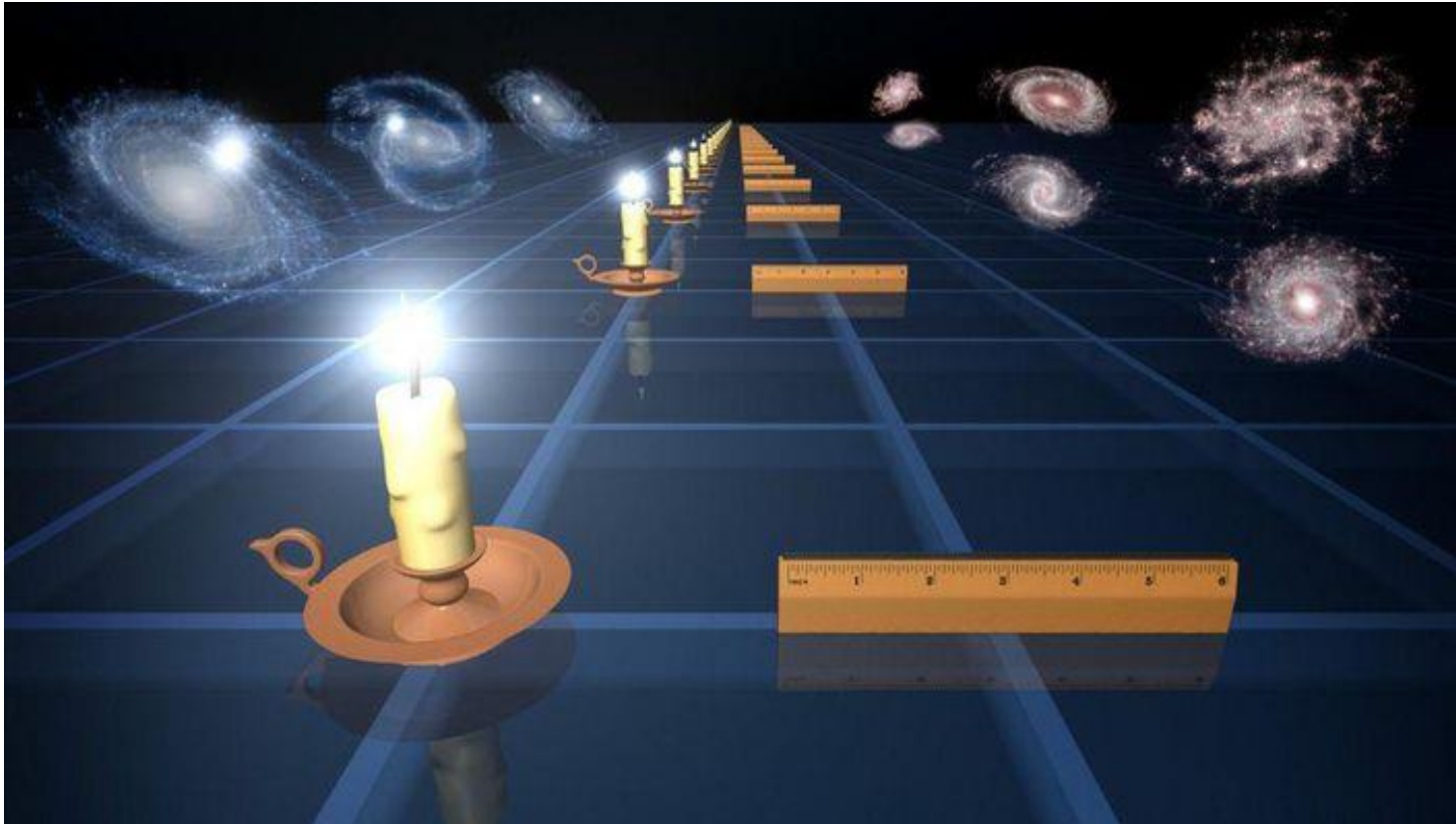
- 3.1 **primordial nucleosynthesis (*BBN*)**
the first three minutes
- 3.2 **cosmic microwave background radiation – essentials:**
formation, black body radiation, fluctuations
- 3.3 **cosmic microwave background radiation – experiments:**
COBE, WMAP & Planck
- 3.4 **cosmological Λ CDM concordance model**

4. Structure Formation in the Universe

- 4.1 Inflation & Early Universe
- 4.2 *BAO* – Baryon Acoustic Oscillations :
formation & relevance in the present universe
- 4.3 Large–Scale Galaxy Surveys:
the matter power spectrum
- 4.4 Evolution of Large–Scale Structures:
Hot, Warm & Cold Dark Matter

5. Dark Universe

- 5.1 **Evidences for Dark Matter:**
from *Vera Rubin & Fritz Zwicky* to the Bullet–Cluster
- 5.2 **Gravitational Lenses for the Dark Universe:**
strong & weak lensing
- 5.3 **Dark Matter Halos: the *NFW* – Profile**
- 5.4 **Dark Energy:**
the future evolution of the Universe



"I'VE SEEN OUT TO THE LIMIT OF THE
OBSERVABLE UNIVERSE, AND BELIEVE ME, IT'S NO
BETTER OUT THERE THAN IT IS HERE."

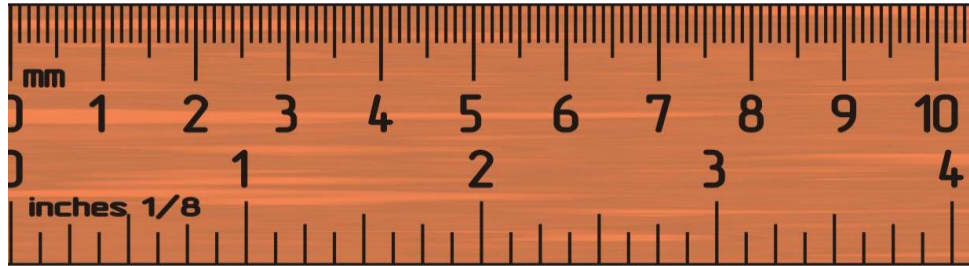
Q: NASA, S. Harris

CHAPTER 1 – INTRODUCTION

cosmological distances and time scales

■ Measurement of distances

- cosmological standard candles up to $D = ?$

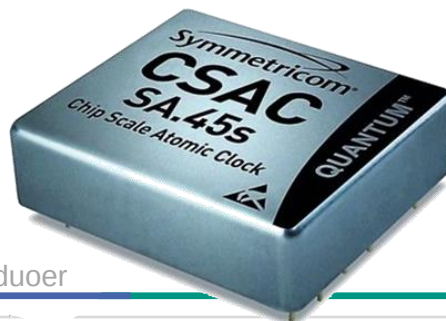


Q: NASA; magns

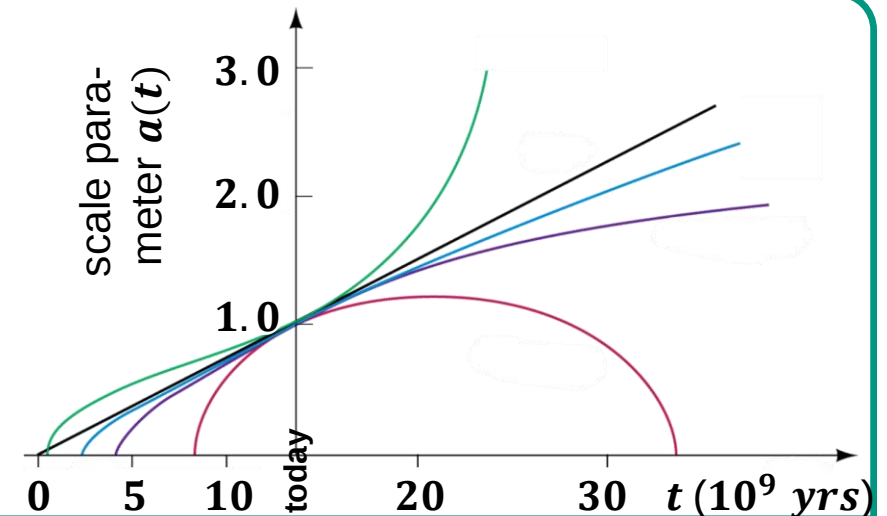


■ Measurement of cosmological time scales

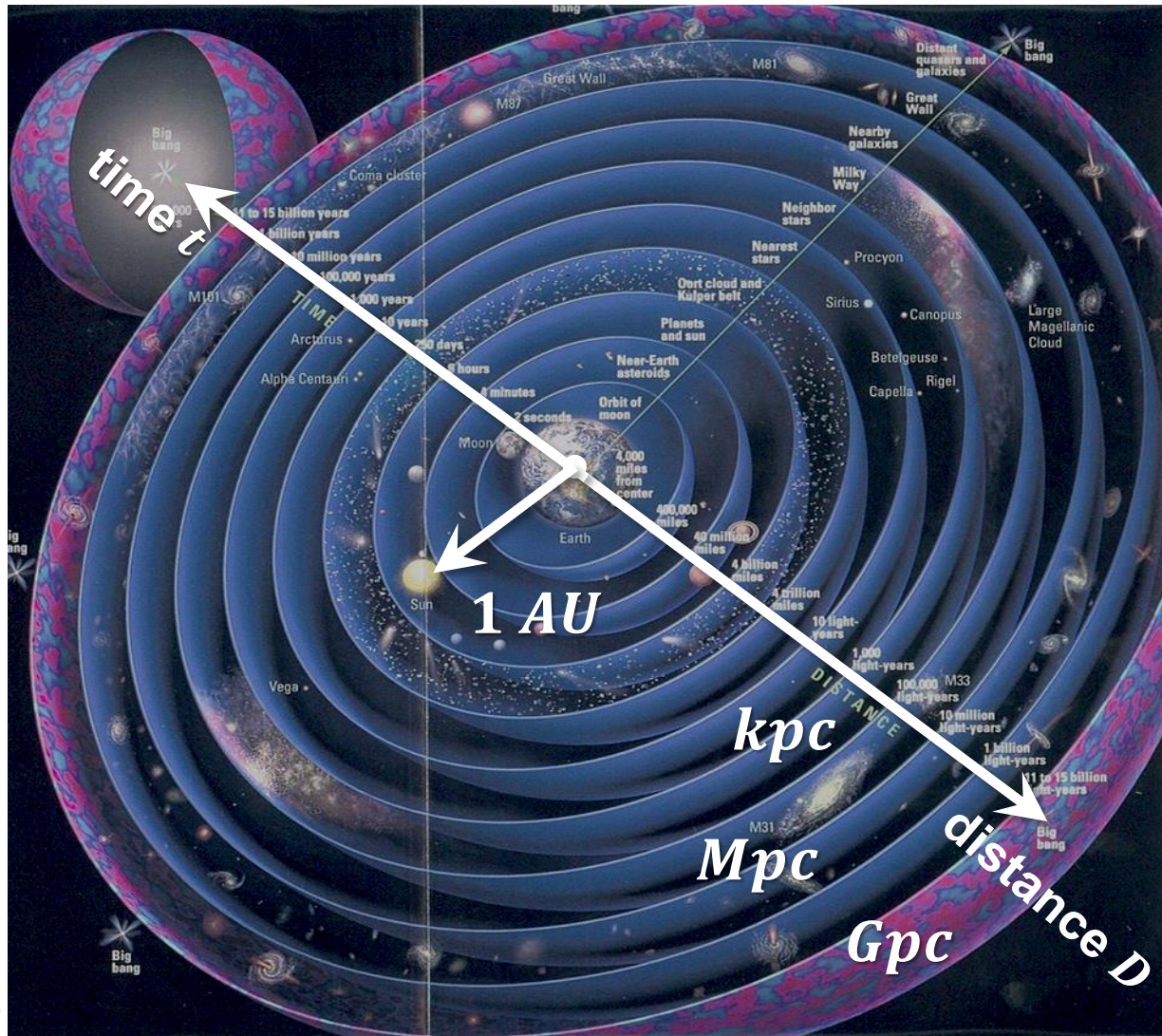
- cosmological standard time scale up to $t = ?$



Q: symmetricom, duoer



Cosmology: typical time– and distance– scales



astronomical scales

$$1 \text{ AU} = 1.496 \times 10^{11} \text{ m}$$

$$1 \text{ ly} = 9.461 \times 10^{15} \text{ m}$$

$$= 63\,240 \text{ AU}$$

$$= 0.3066 \text{ pc}$$

$$1 \text{ pc} = 3.086 \times 10^{16} \text{ m}$$

$$= 2.06 \times 10^5 \text{ AU}$$

$$= 3.262 \text{ ly}$$

AU: Astronomical Unit

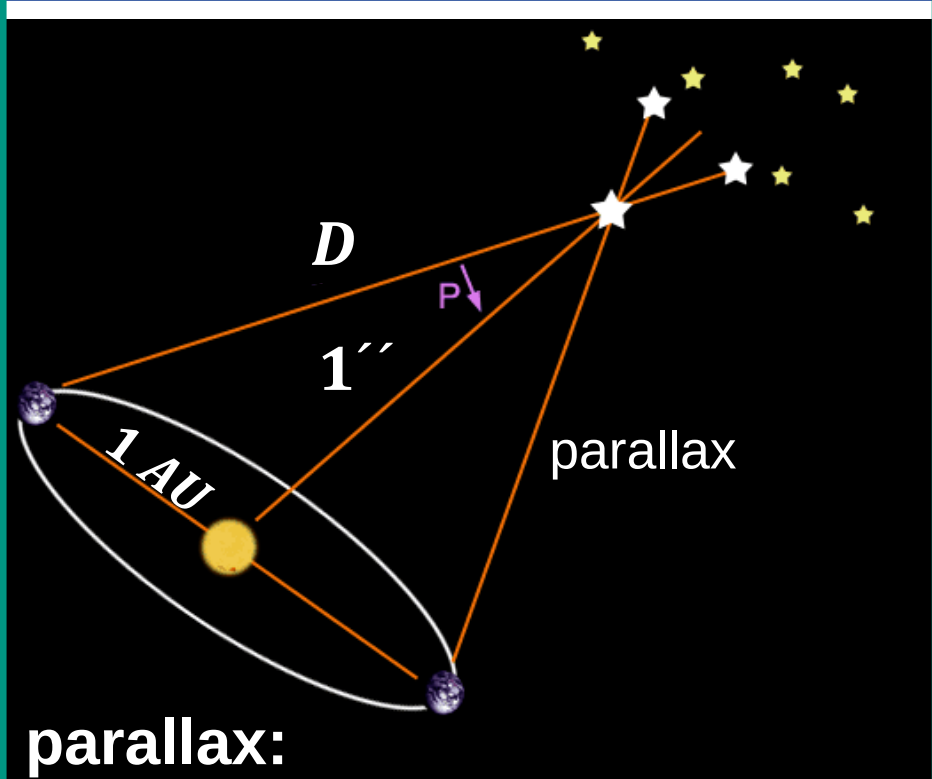
major semi–axis of Earth around Sun

ly: light–year light trajectory in $t = 1 \text{ yr}$

pc: parsec parallax of $1 \text{ arc} - \text{sec}$

Cosmology: distance – scale pc

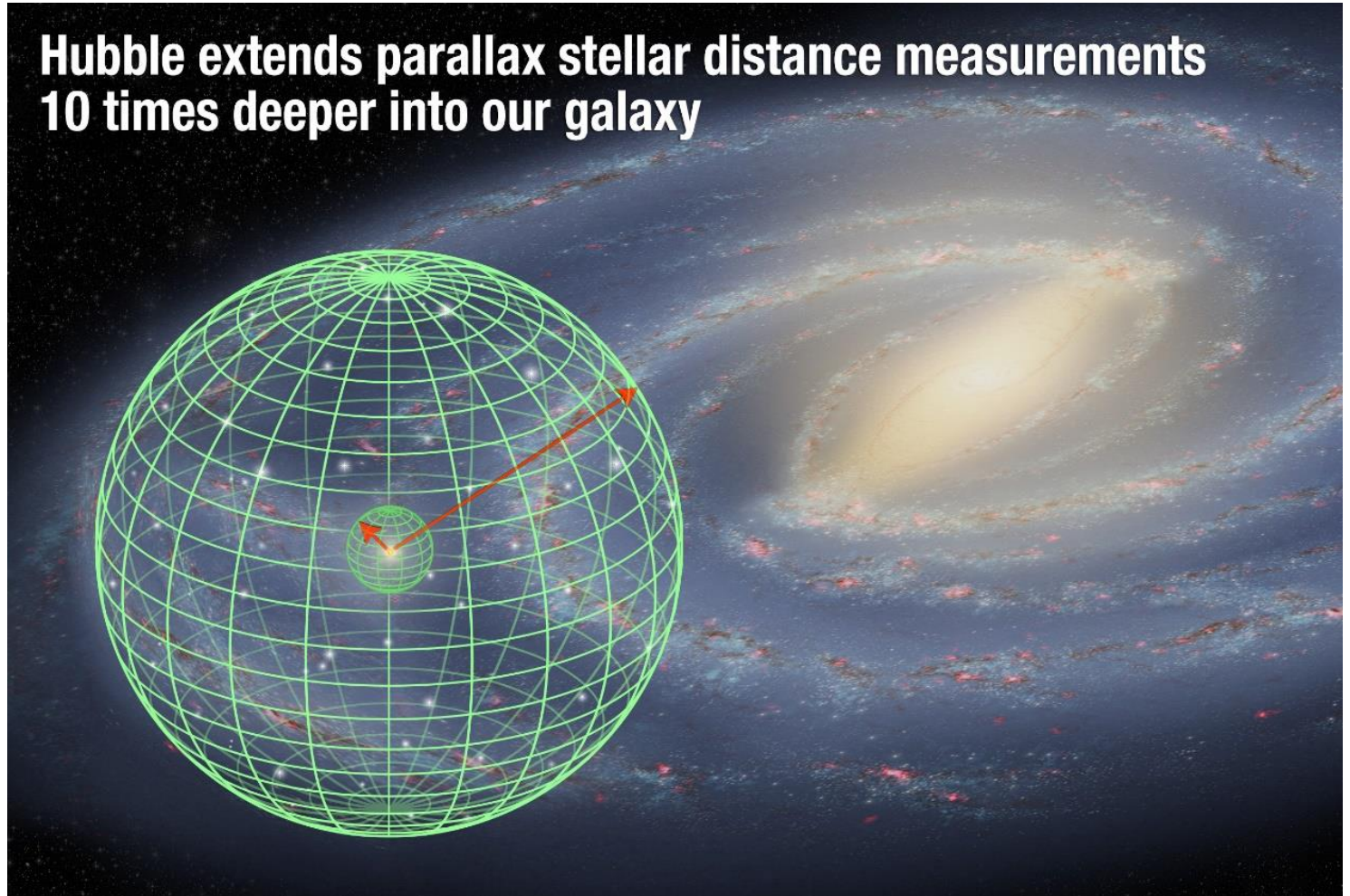
definition of 'parallax'



parallax:

$$\begin{aligned} 1 \text{ pc} &= 1 \text{ AU subtends to } 1'' \\ &= 1 \text{ AU} \cdot 180 \cdot (3600 / \pi) \end{aligned}$$

**Hubble extends parallax stellar distance measurements
10 times deeper into our galaxy**



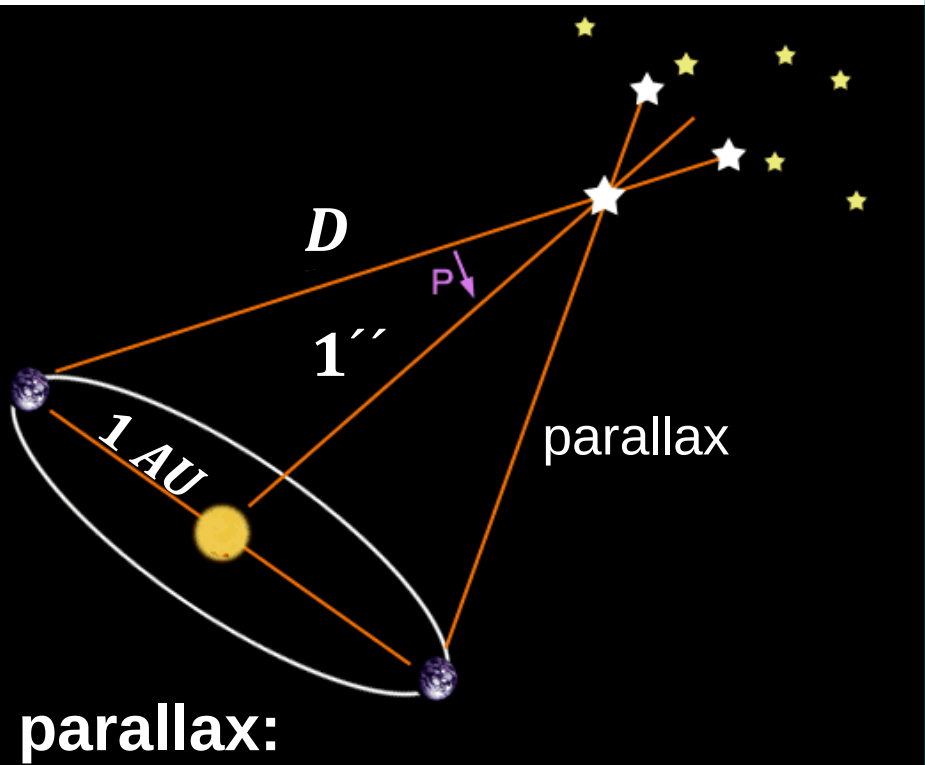
Q: NASA

2014 – parallax measurements now up to 3 – 4 kpc

$$1'' = 1 \text{ arcsec} = 1 \text{ arcsecond}$$

Cosmology: distances & *GAIA* observatory (*ESA*)

definition of 'parallax'



parallax:

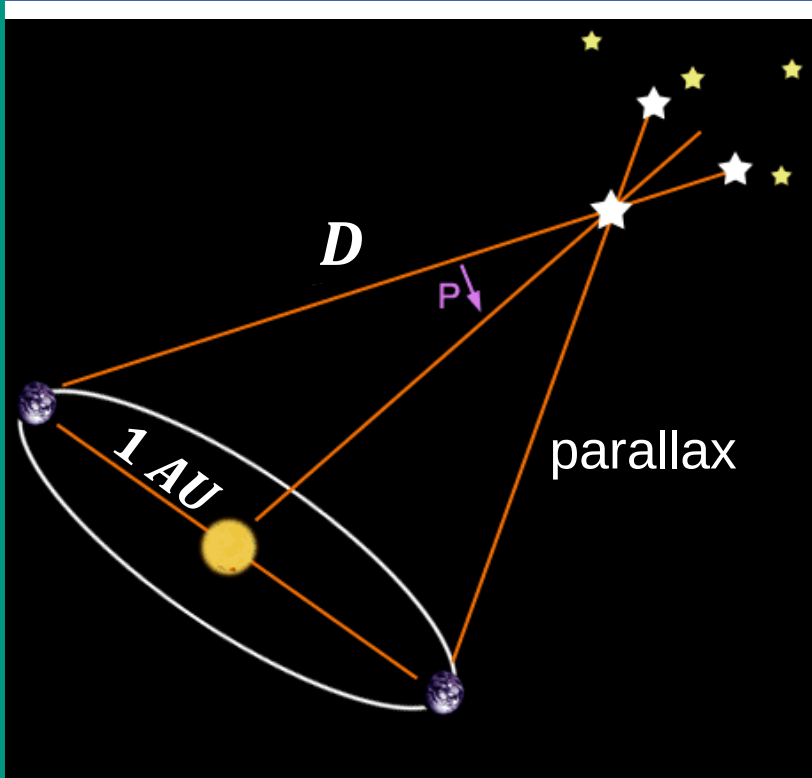
$$\begin{aligned} 1 \text{ pc} &= 1 \text{ AU subtends to } 1'' \\ &= 1 \text{ AU} \cdot 180 \cdot (3600 / \pi) \end{aligned}$$



measurements of *GAIA* with μas – precision*

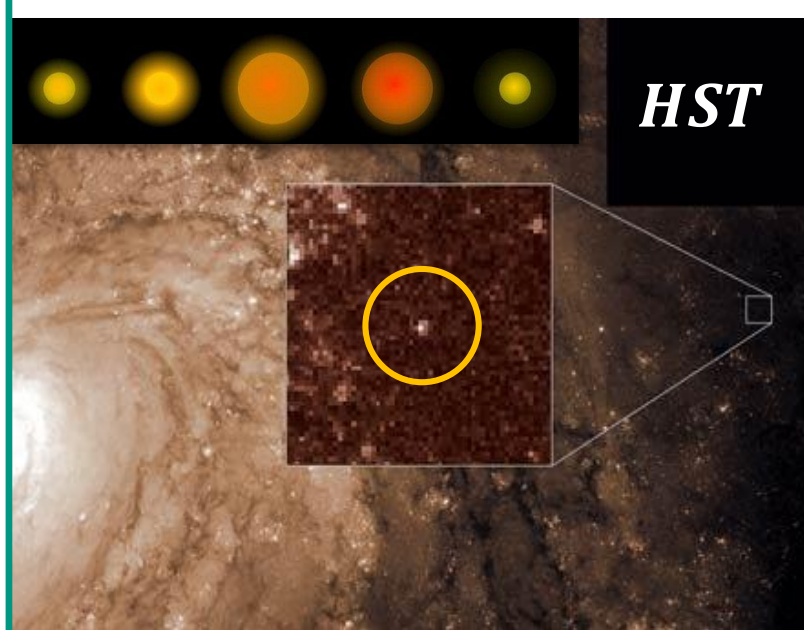
Cosmology: the 'classical' distance ladder

parallax method



Parallax method: reach to
 $D = \text{few kpc}$ (Milky Way)

δ – Cepheids



δ – Cepheids: absolute
luminosity $L_{abs} \sim 10^4 L_{\odot}$

$D = 1 \text{ kpc} \dots 50 \text{ Mpc}$

$SN - Ia$ – brightness



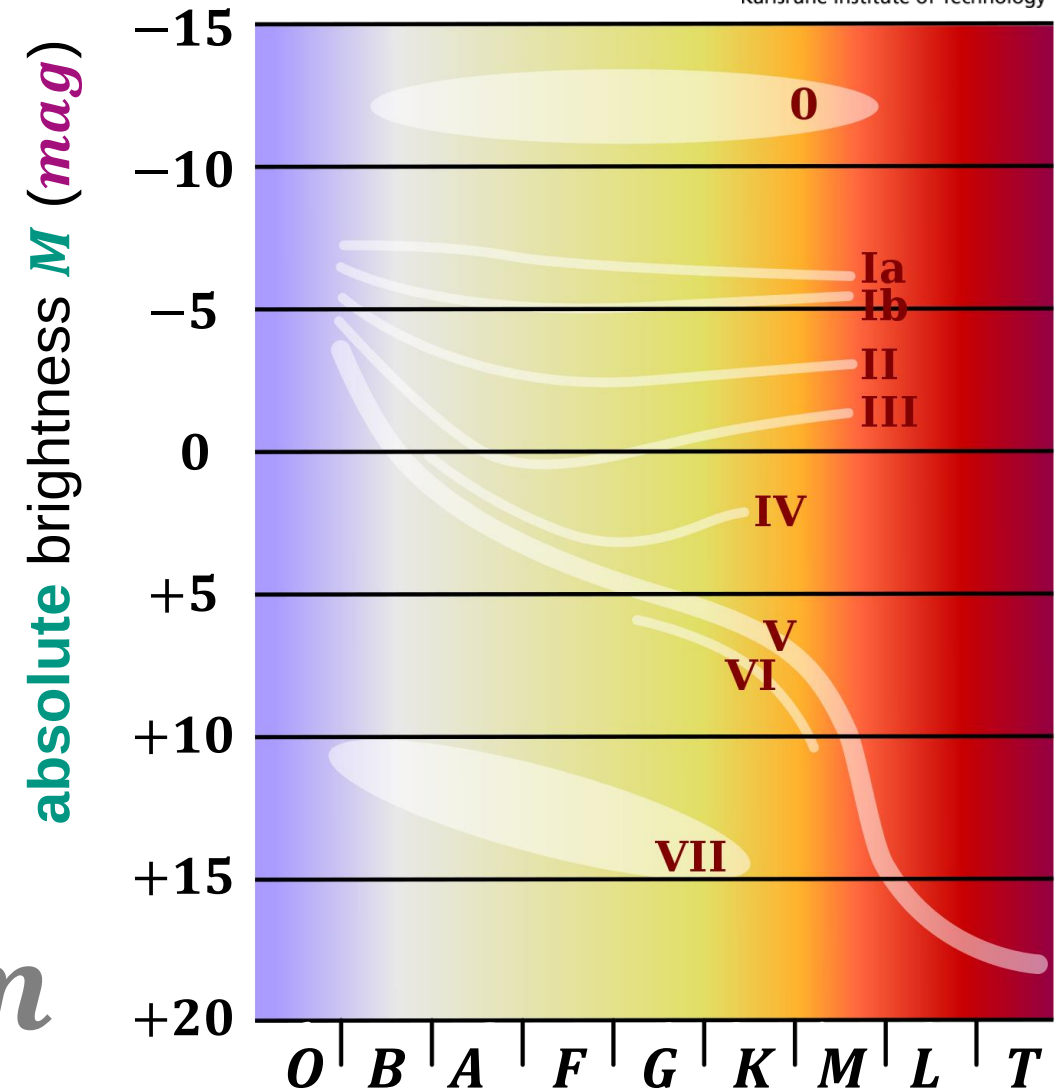
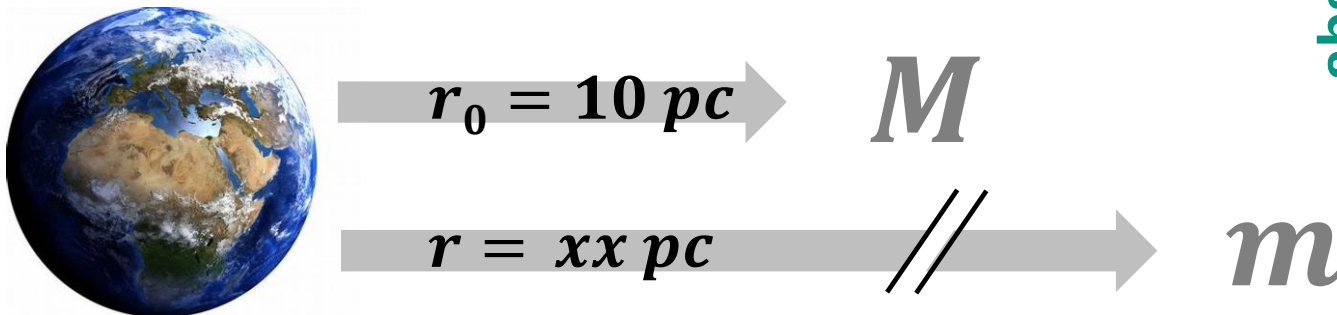
Supernovae Type Ia:
luminosity $L_{abs} \sim 10^9 L_{\odot}$

$D = 30 \text{ Mpc} \dots 3 \text{ Gpc}$

Brightness – definition in astronomy

■ absolute & relative brightness

- brightness of objects is measured in **magnitudes (*mag*)** = logarithmic scale
- **absolute** brightness ***M***:
defined at unit distance $r_0 = 10 \text{ pc}$
- **apparent** brightness ***m***:
object as it appears **here on Earth**



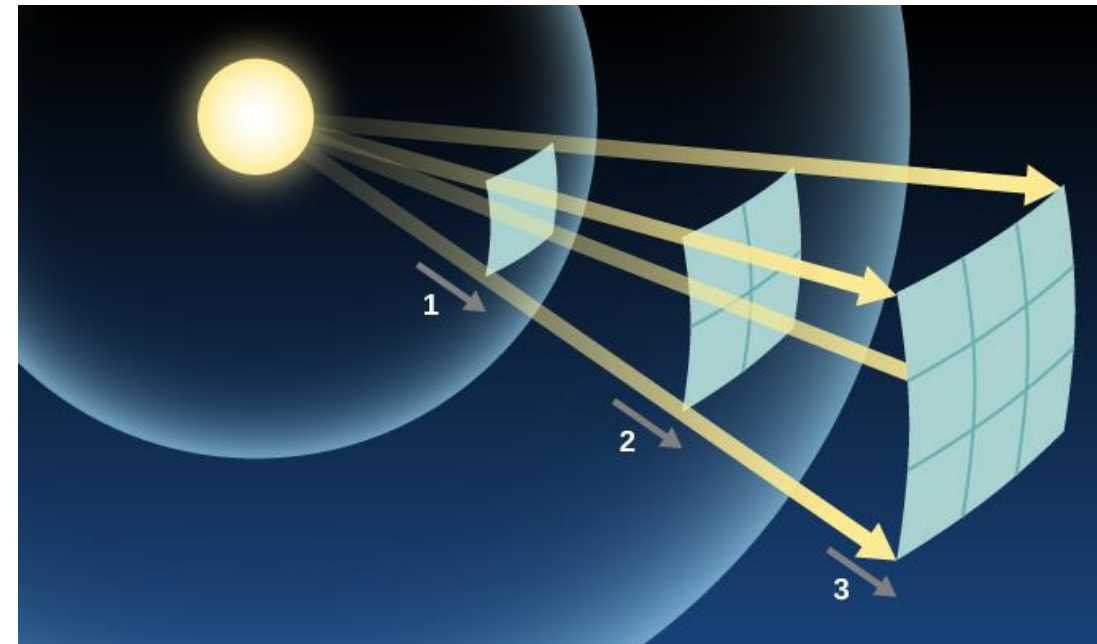
Brightness – definition in astronomy

■ absolute & apparent brightness: **magnitude**

- reference: **$m \equiv 0$** is defined for the nearby star **Wega** (α *Lyrae*)
- **mag** = logarithmic unit $\Leftrightarrow$ difference of **$5 mag \equiv factor 100$** in brightness
per magnitude: **$factor = \sqrt[5]{100} = 2.512$**
- **Distance modulus $m - M$:**

$$m - M = 5 \cdot \log \left(\frac{r}{r_0} \right)$$

quadratic
distance law

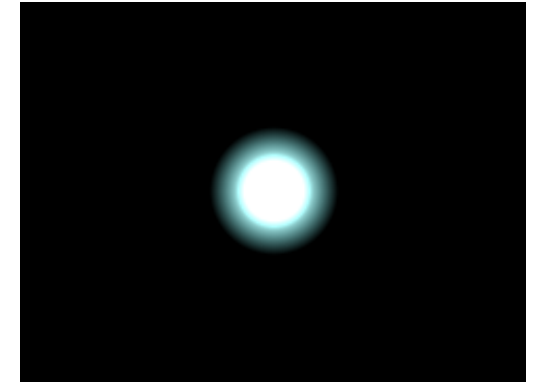


Cepheids – cosmological standard candles #1

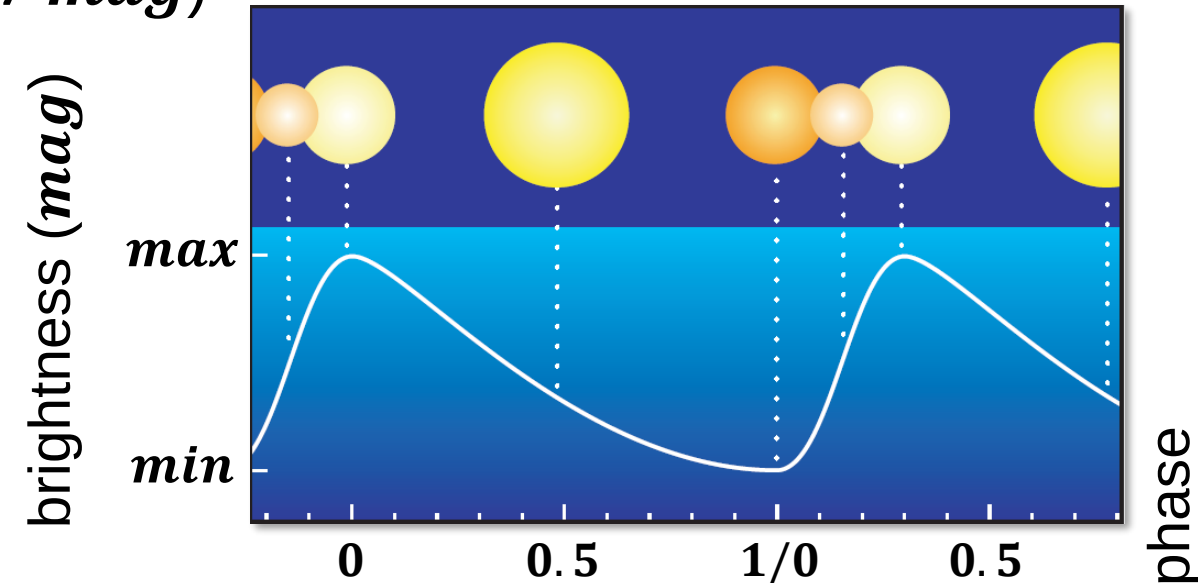
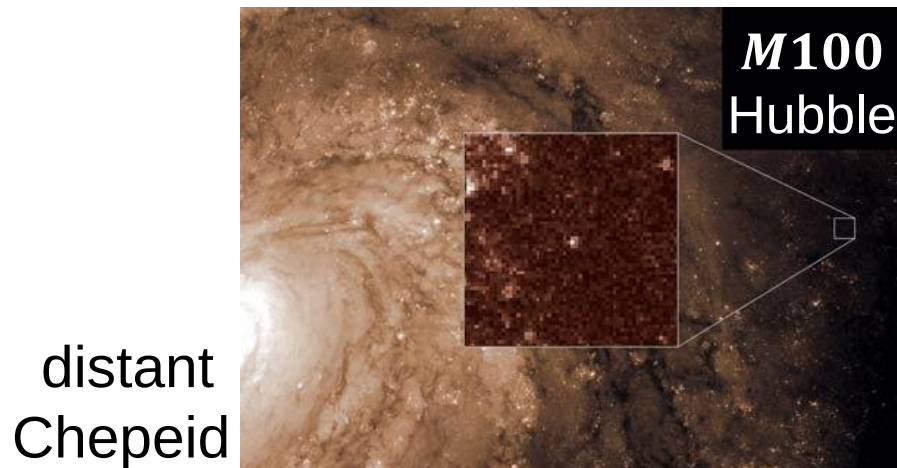
■ Cepheids:

- giant stars with radial pulsations (‘super–giants’)
- allows to determine **absolute brightness from period T_0**
- mass range: $M = 5 \dots 15 M_{\odot}$ ($\Rightarrow$ very rare stars)
- spectral class: F, G (white, yellow)
- brightness class*: Ia, II ($M_V = -2 \dots -7 \text{ mag}$)
- period: $T_0 < 100$ days

period of a δ – Cepheid



Q: NASA



M_V = visual brightness

Cepheids – cosmological standard candles #1

■ ‘Kappa mechanism’: the underlying cause of radial pulsation

changes in the **opacity** $\kappa(p, T)$ of a specific stellar layer

- key: hot stellar core with an **He – ionisation layer** of He^{++} & He^+ – ions

- **cycle of stellar pulsation:**

increasing T :

more He^{++} & free electrons

⇒ larger opacity

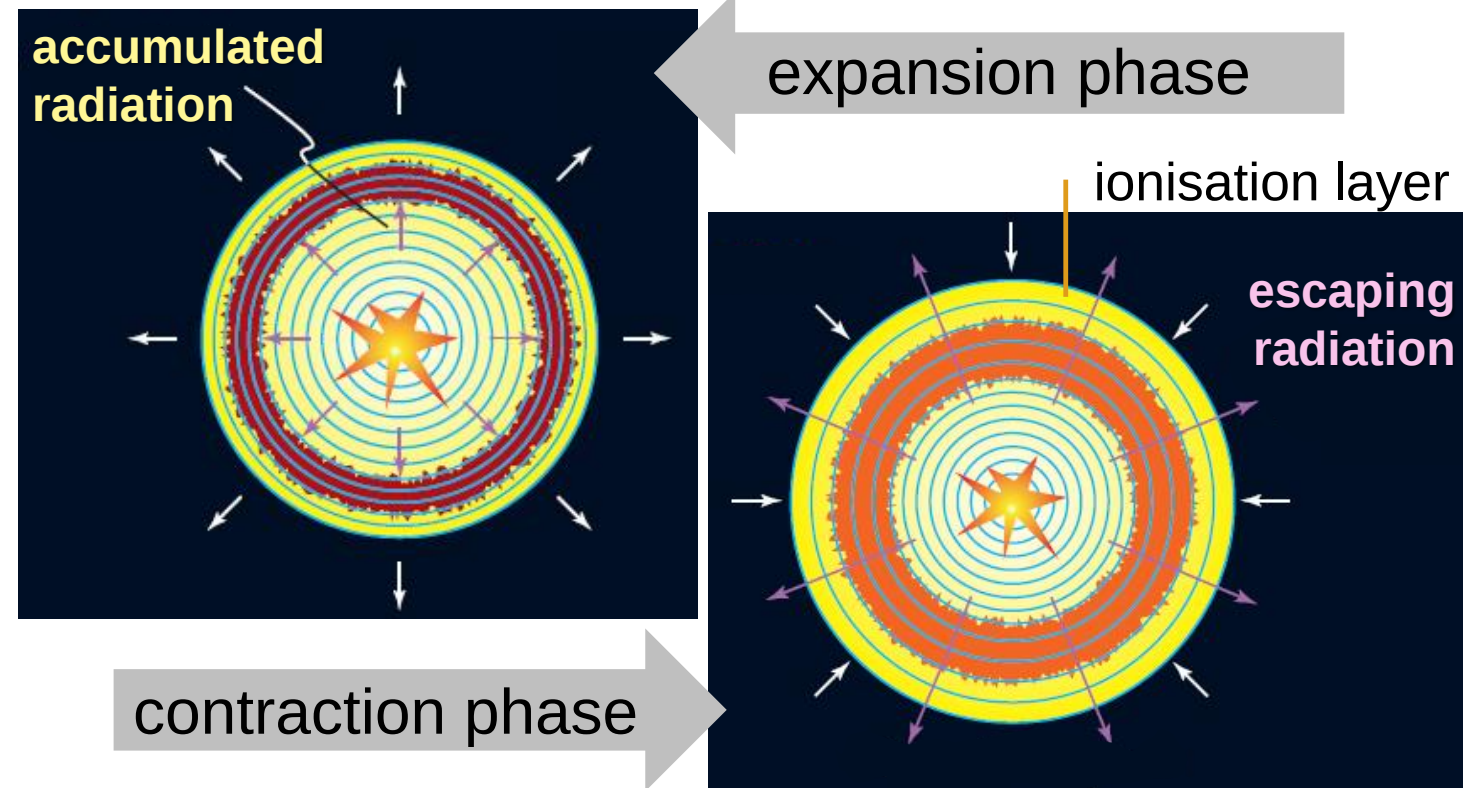
⇒ increase of rad. pressure

⇒ expansion of outer shell

decreasing T :

gravitational contraction

⇒ recombination to He^+



Cepheids – cosmological standard candles #1

■ pulsation time scale of outer stellar shell

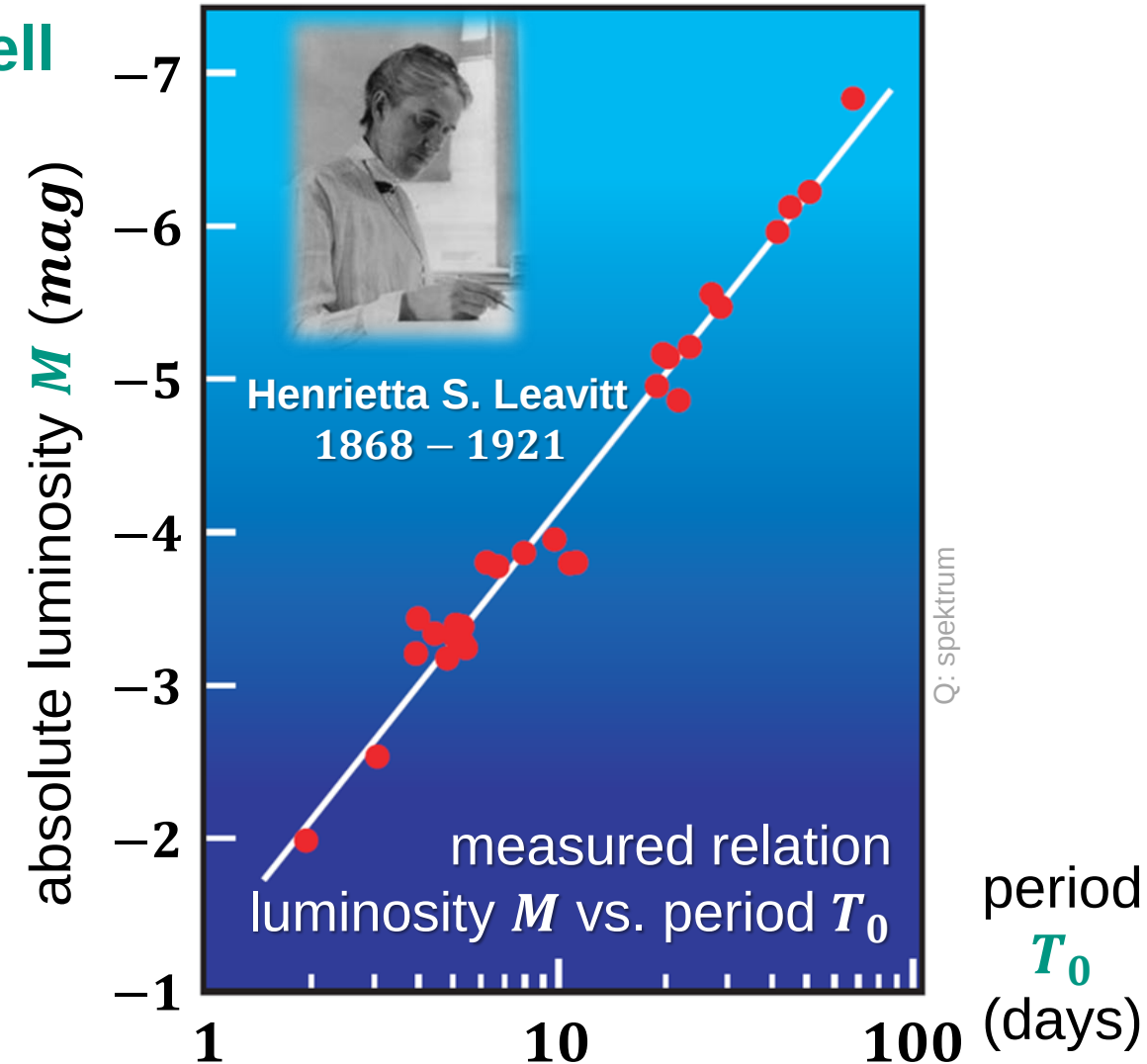
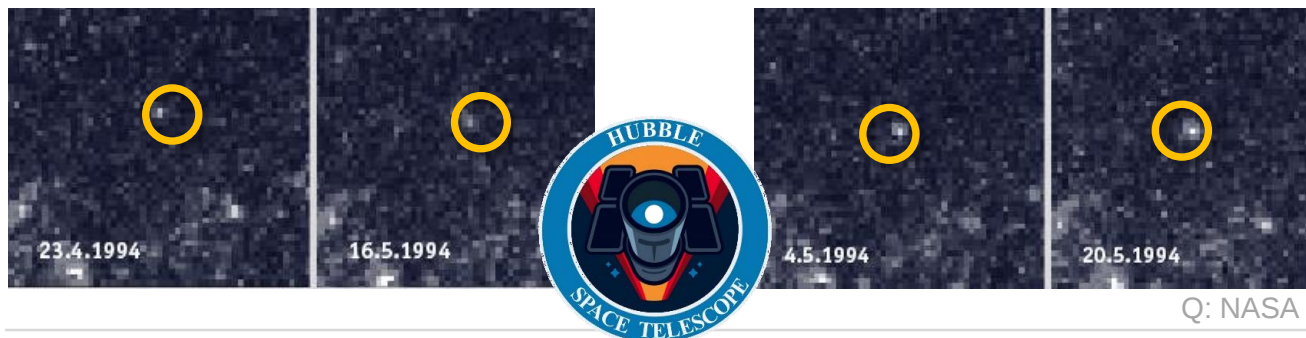
- very massive, very bright cepheids are pulsating more slowly, **period** $T_0 \sim L^{3/4}$

pulsation period – luminosity relation

$$M = -2.81 \cdot \log(T_0) - (1.43 \pm 0.1)$$

T_0 : pulsation period (*days*)

M : absolute brightness (*mag*)



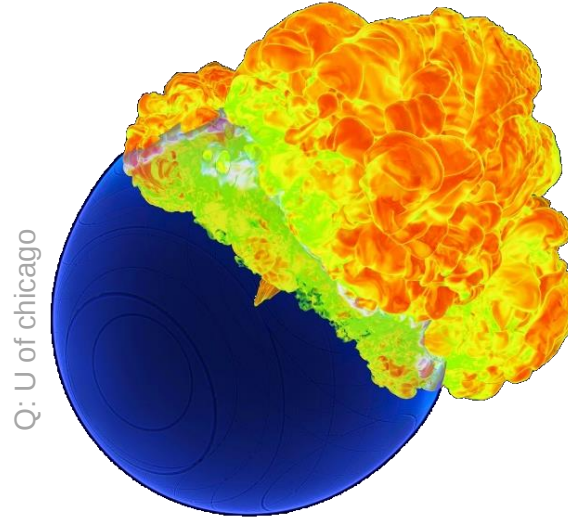
Supernovae – cosmological standard candles #2

■ *SN Ia*: thermonuclear detonation*

- a White Dwarf reaches a mass of $M \geq 1.4 M_{\odot}$



merger scenario



accretion scenario

